

Predictions of the mutual phenomena of the Galilean Satellites of Jupiter for 1985–1986

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Summary. In 1985 and 1986, the Earth and the Sun will be in the equatorial plane of Jupiter, so that mutual phenomena (occultations and eclipses) will occur between the Galilean Satellites. The present paper gives the calculated dates, the durations and the magnitudes of each observable phenomenon for three different ephemerides of the Galilean Satellites. The instrumental possibilities of observations are given and the significance of those observations is explained.

Key words: mutual phenomena – galilean satellites

Introduction

The configuration of the orbits of the Galilean Satellites of Jupiter induces, two times each jovian year of 11.6 yr, phenomena between the satellites themselves. The four satellites have orbits which are nearly in the equatorial plane of Jupiter. When the Earth goes through this plane, i.e. when the jovian declination of the Earth becomes zero, the satellites may occult one another for a terrestrial observer. Similarly, when the Sun goes through the equatorial plane of Jupiter, i.e. when the jovian declination of the Sun becomes zero, the satellites may enter the umbra or the penumbra of the other satellites. Because of the small size of the satellites and the very small inclination of their orbits on the jovian equator, mutual phenomena do not occur for each geocentric conjunction (for the occultations) or heliocentric one (for the eclipses) during the favorable period. This favorable period occurs when the jovicentric declinations of the Earth and the Sun are smaller than a defined quantity. These phenomena are easily predictable with modern computers and their observation – which presents no major difficulties – gives interesting information about the Galilean Satellites themselves. In 1985, the most favorable period for the phenomena will occur during the opposition of Jupiter with the Sun so that the observations will be easier. Figure 1 shows the variations of the jovicentric declinations of Earth and Sun for 1979, 1985, and 1992 periods. We may see that the 1992 period is more favorable than the 1979 one but less than the 1985 one. Further, the value of the equatorial geocentric declination of Jupiter is also important to determine the observational conditions: the observatories are more numerous in the northern than in the southern hemisphere.

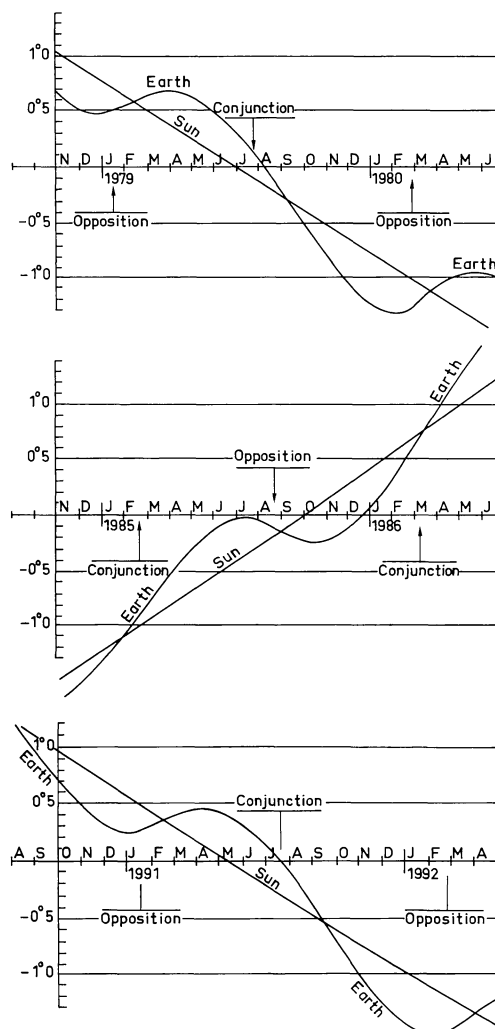


Fig. 1. Jovicentric declination of the Earth and the Sun for the periods of 1979, 1985, and 1991

The predictions for 1985 and 1986

For the calculations, we used the same algorithm as for the previous period (Arlot, 1978). This algorithm gives a precision which varies with each phenomenon: the faster the relative

Table 1. Dates and elements of the mutual phenomena of 1985-86 for 3 ephemerides, G-5, E-2, and SV2

Table with columns: Year m. day Phenomenon, Dates of the beginning (5), Date of the maximum: G-5, E-2, SV2, Duration (6): G-5, E-2, SV2, Magnitude: G-5, E-2, SV2, Impact parameter (11): G-5, E-2, SV2, DJ (Rj) Notes: (19) (20). The table contains multiple rows of data for various dates in 1985 and 1986, detailing astronomical observations and calculations.

Table 1 (continued)

Table with columns: Year, m., day, Phenomenon, Dates of the beginning (h m s), Date of the maximum (G-5, E-2, SV2), Duration (G-5, E-2, SV2), Magnitude (G-5, E-2, SV2), Impact parameter (G-5, E-2, SV2), DJ (Rj), Notes.

Table 1 (continued)

Year	m.	day	Phenomenon	Dates of the beginning						Date of the maximum:			Duration ⁽⁵⁾ :			Magnitude :			Impact parameter ⁽¹¹⁾			Dj (R _J)	Notes:											
				(5)		(6)		(7)		G-5	E-2	SV2	G-5	E-2	SV2	G-5	E-2	SV2	G-5	E-2	SV2													
				h	m	s	h	m	s	h	m	s	m	s	m	s	m	s	m	s	m			s										
1985	11	29	3 OCC 1 P		4	49	30.		4	54	57.		55	8.	55	14.		659	658	656		0.543	0.539	0.541		0.639	0.643	0.642		5.9				
1985	11	29	3 OCC 1 P		16	13	4.		16	19	54.		19	20.	22	30.		821	811	1002		0.233	0.227	0.369		0.931	0.937	0.803		0.3	(3)			
1985	11	30	3 ECL 1 P	6	45	22.		6	47	54.		6	50	6.				267	268	264		14.3%	14.6%	13.8%		0.634	0.631	0.640		4.8				
1985	11	30	3 ECL 2 P	10	49	55.		10	52	46.		10	54	26.		54	21.	54	0.	198	196	214		9.5%	9.2%	12.3%		0.608	0.611	0.577		6.2		
1985	12	1	1 OCC 2 P		4	5	38.		4	7	11.		7	2.	7	44.		185	192	240		0.210	0.229	0.389		0.699	0.683	0.557		4.8				
1985	12	2	2 ECL 1 P	12	40	53.		12	42	30.		12	43	20.		43	55.		155	153	161		12.4%	12.1%	12.5%		0.699	0.683	0.557		4.8			
1985	12	4	1 OCC 2 P		17	17	30.		17	19	14.		19	5.	19	48.		207	213	250		0.288	0.308	0.461		0.631	0.615	0.495		3.7				
1985	12	6	2 ECL 1 P	1	48	14.		1	49	28.		1	50	40.		50	35.		148	145	154		10.6%	10.1%	11.8%		0.367	0.378	0.337		4.0			
1985	12	6	3 OCC 1 P		14	56	57.		14	58	14.		14	58	14.		58	25.		2164	2178	2157		0.715	0.711	0.757		0.468	0.472	0.429		5.4		
1985	12	6	3 OCC 1 P		13	25	36.		13	25	36.		13	25	36.		30	26		2345	2355	2419		0.591	0.586	0.685		0.584	0.588	0.496		3.3		
1985	12	7	3 ECL 1 P	9	50	12.		9	53	42.		9	53	42.		53	47	52.						0.591	0.586	0.685		0.584	0.588	0.496		3.3		
1985	12	7	3 ECL 2 P	14	10	46.		14	13	10.		14	15	30.		15	30.		282	282	289		41.6%	41.4%	47.0%		0.930	0.927	0.919		4.2	(2)		
1985	12	7	4 OCC 1 P		16	38	44.		16	41	3.		16	43	4.		43	55.		278	279	278		0.491	0.495	0.489		0.617	0.613	0.619		5.2	(3)	
1985	12	7	1 OCC 2 P		6	29	38.		6	31	31.		6	31	31.		31	22.		226	231	259		0.376	0.396	0.543		0.557	0.542	0.428		4.6		
1985	12	9	2 OCC 3 P	14	55	39.		5	5	14.		5	5	14.		5	5	14.		718	715	782		0.487	0.480	0.658		0.623	0.632	0.397		13.0		
1985	12	9	2 OCC 3 P		1	7	57.		1	7	57.		1	9	36.		9	36.		199	206	144		7.9%	7.6%	9.6%		0.423	0.429	0.397		13.0		
1985	12	10	1 OCC 3 P		8	19	59.		8	21	43.		8	21	43.		21	42.		207	205	192		0.197	0.197	0.232		0.895	0.883	0.741		4.9		
1985	12	11	1 OCC 2 P	4	3	9.		4	4	31.		4	5	30.		4	5	30.		124	121	134		5.7%	5.2%	7.2%		0.466	0.477	0.438		4.2	(3)	
1985	12	13	3 OCC 1 P		10	10	15.		10	11	33.		10	13	33.		11	11		157	150	184		0.068	0.062	0.094		1.048	1.054	1.025		5.1		
1985	12	14	3 OCC 1 P		12	48	13.		12	50	18.		12	51	18.		50	23.		50	33.			0.0	0.0	0.0		1.221	1.219	1.195		3.6	(2)	
1985	12	14	3 ECL 2 A	17	30	52.		17	33	8.		17	35	35.		17	35	35.		298	298	298		0.0	0.0	0.0		1.221	1.219	1.195		3.6	(2)	
1985	12	15	1 OCC 2 P		8	53	55.		8	56	1.		8	56	1.		55	53.		253	256	271		0.578	0.599	0.732		0.952	0.948	0.869		11.6		
1985	12	15	1 OCC 4 A		20	4	54.		20	8	27.		20	8	27.		8	14.		427	427	427		0.880	0.874	0.849		0.001	0.008	0.038		4.2		
1985	12	16	2 ECL 1 P	17	10	44.		17	12	11.		17	13	7.		13	2.		109	104	120		3.5%	3.0%	4.8%		0.513	0.523	0.484		4.2			
1985	12	16	2 ECL 1 P		20	9	57.		20	9	57.		20	9	57.		20	9	57.		2521	2549	2505		0.543	0.541	0.699		0.401	0.404	0.216		5.8	
1985	12	17	2 OCC 3 P		4	37	57.		4	40	52.		4	42	40.		42	40.		349	352	373		0.435	0.445	0.552		0.466	0.453	0.314		5.3		
1985	12	17	1 OCC 3 P		11	15	2.		11	17	26.		11	17	26.		17	26.		288	287	295		0.433	0.436	0.529		0.539	0.543	0.583		1.8		
1985	12	18	1 OCC 2 P		22	6	5.		22	8	16.		22	8	16.		7	8.		262	264	274		0.692	0.713	0.840		0.305	0.289	0.192		3.3		
1985	12	20	2 ECL 1 P	6	18	25.		6	19	58.		6	20	42.		20	37.	21	18.		88	81	101		1.6%	1.2%	2.7%		0.556	0.567	0.528		4.9	(2)
1985	12	21	3 OCC 1 P		13	22	35.		13	25	3.		13	25	3.		25	5.		296	294	297		0.358	0.353	0.362		0.773	0.777	0.770		4.5		
1985	12	21	3 ECL 2 P	20	50	40.		20	53	1.		20	55	10.		20	55	10.		262	262	251		42.4%	42.2%	34.8%		0.279	0.281	0.335		5.3		
1985	12	22	1 OCC 2 P		11	18	24.		11	20	37.		11	20	37.		20	29.		21	14.			267	269	274		0.815	0.836	0.956		4.2		
1985	12	23	2 ECL 1 P	19	26	8.		19	27	57.		19	28	20.		28	20	28	55.		53	33		0.0	0.0	0.0		0.210	0.194	0.103		4.2	(2)	
1985	12	24	4 OCC 1 P		1	8	16.		1	18	51.		1	18	51.		18	23.		312	312	1301		0.627	0.618	0.593		0.598	0.609	0.571		4.4	(2)	
1985	12	24	4 OCC 1 P		6	3	31.		6	17	18.		6	17	18.		17	42.		1595	1588	1520		0.770	0.765	0.645		0.474	0.482	0.504		5.8		
1985	12	24	2 OCC 3 A		8	11	5.		8	14	26.		8	14	26.		14	26.		402	402	399		0.769	0.779	0.701		0.031	0.018	0.119		5.7		
1985	12	24	2 OCC 3 P		14	12	46.		14	15	34.		14	15	34.		15	34.		336	336	333		0.680	0.678	0.646		0.215	0.217	0.259		2.6		
1985	12	25	2 OCC 3 P		2	38	49.		2	38	49.		2	38	49.		38	35.		433	433	425		1.018	1.019	0.907		0.125	0.125	0.224		4.3	(13)	
1985	12	26	1 OCC 2 P		0	30	41.		0	32	56.		0	32	56.		36	5.		269	270	272		0.946	0.968	1.081		0.110	0.094	0.009		4.1		
1985	12	27	2 ECL 1 P	8	33	59.		8	36	10.		8	36	10.		36	5.		36	6.			0.0	0.0	0.0		0.110	0.094	0.009		4.1			
1985	12	28	3 OCC 1 P		16	29	38.		16	32	31.		16	32	31.		32	33.		345	344	342		0.708	0.703	0.691		0.454	0.457	0.469		4.0	(2)	
1985	12	29	3 OCC 1 P	0	9	30.		0	12	21.		0	13	36.		13	31.		143	141	88		5.4%	5.2%			0.582	0.585	0.646		5.0			
1985	12	30	1 OCC 2 T		13	43	5.		13	45	19.		13	45	19.		45	11.		267	267	265		1.086	1.079	0.972		0.006	0.011	0.090		3.9	(14)	
1985	12	31	1 OCC 2 P		21																													

Table 1 (continued)

Year	m.	day	Phenomenon {1} {2} {3} {4}	Dates of the beginning						Date of the maximum:			Duration ⁽⁵⁾ :			Magnitude :			Impact parameter ⁽¹¹⁾			D _J (R _J)	Notes: {19} {20}										
				{5}		{6}		{7}		{8}		{9}		{10}		{11}		{12}		{13}				{14}		{15}		{16}		{17}		{18}	
				h	m	s	h	m	s	h	m	s	m	s	m	s	m	s	m	s	%			%	%	%	%	%	%	%	%	%	
1986	4	2	2 ECL 1	4	30	32.	4	34	15.	34	15.	34	45.	0	%	0	%	0	%	0.759	0.768	0.824	5.9	(2)									
1986	4	4	2 ECL 3 P	14	15	56.	14	22	51.	14	49	49.	48	28.	50	50.	3379	3411	3579	16.3%	17.0%	18.9%	0.450	0.432	0.356	2.6							
1986	4	4	2 ECL 3 A	22	20	55.	22	29	32.	22	55	48.	56	34.	56	39.	2970	2970	3023	14.9%	14.9%	14.9%	0.219	0.204	0.148	6.7							
1986	4	5	2 OCC 3 P				3	59	28.	4	8	44.	9	2.	9	47.	1103	1114	1175	0.234	0.240	0.274	0.704	0.696	0.653	9.2							
1986	4	5	2 ECL 1	17	52	30.				17	57	18.	57	18.	57	39.				0	%	0	%	0.713	0.722	0.787	5.9						
1986	4	5	2 OCC 1 P				19	56	35.	19	58	45.	58	47.	58	57.	262	180	0	0.023	0.011	0.0	0.788	0.798	0.899	5.6							
1986	4	6	2 OCC 1 P				3	27	30.	3	34	51.	34	50.	37	12.	883	819	0	0.113	0.096	0.0	0.709	0.724	0.914	0.7							
1986	4	6	2 ECL 1	5	47	52.				5	53	12.	53	7.	55	9.				0	%	0	%	0.832	0.847	1.197	1.2						
1986	4	9	2 ECL 1 P	7	16	42.	7	20	8.	7	22	30.	22	30.	22	55.	287	263	0	1.4%	1.1%	0	0.663	0.672	0.747	5.9							
1986	4	9	2 OCC 1 P				9	50	16.	9	56	38.	56	42.	56	23.	785	733	0	0.092	0.080	0.0	0.731	0.742	0.862	5.3							
1986	4	9	2 OCC 1 P				14	43	41.	14	56	11.	56	7.	58	51.	1483	1412	0	0.203	0.187	0.0	0.634	0.648	0.830	2.1							
1986	4	9	2 ECL 1 P	17	17	46.	17	21	35.	17	28	4.	27	59.	30	10.	794	728	0	3.4%	2.5%	0	0.671	0.686	1.008	0.1							
1986	4	12	2 ECL 3 P	4	39	33.	4	44	23.	4	53	46.	54	6.	54	37.	1111	1122	1169	9.8%	10.2%	11.4%	0.513	0.502	0.459	7.4							
1986	4	12	2 ECL 1 P	20	47	8.	20	50	16.	20	54	45.	54	45.	55	6.	540	522	261	5.3%	4.8%	0.6%	0.601	0.610	0.697	5.8							
1986	4	13	2 ECL 1 P	4	52	6.	4	54	55.	5	4	59.	4	54.	7	11.	1192	1161	362	13.0%	11.8%	0.2%	0.541	0.555	0.750	1.3							
1986	4	16	2 ECL 1 P	10	24	41.	10	28	5.	10	35	22.	35	27.	35	27.	899	881	620	11.1%	10.4%	3.6%	0.531	0.540	0.642	5.7							
1986	4	16	2 ECL 1 P	16	24	49.	16	27	53.	16	41	6.	40	56.	43	23.	1552	1529	1015	21.9%	20.7%	5.5%	0.443	0.458	0.643	2.4							
1986	4	19	2 ECL 3 P	9	22	5.				9	30	17.	30	27.	31	7.	75	277	0	0	%	0.0%	0.4%	0.862	0.854	0.819	7.5						
1986	4	20	2 ECL 1 P	0	24	39.	0	29	22.	0	46	2.	46	17.	45	22.	2275	2254	1750	20.5%	19.6%	9.2%	0.435	0.445	0.571	5.4							
1986	4	20	2 ECL 1 P	3	16	1.	3	23	3.	3	49	36.	49	36.	45	33.	2885	2861	2218	28.8%	27.6%	12.1%	0.369	0.381	0.553	3.8							
1986	4	26	2 ECL 3							13	37	10.	37	19.	40	50.	0	0	0	0	%	0	%	1.381	1.375	1.213	3.8						

Notes to Table 1

- (1) The apparent distance between the Sun and Jupiter is smaller than 25° but larger than 10°.
- (2) Grazing phenomena, the magnitude is less than 0.03 for occultations and less than 2% for eclipses.
- (3) The distance between the eclipsed or occulted satellite and the limb of Jupiter is less than 0.3 jovian radius (the event may occur in front of Jupiter).
- (4) Double phenomenon, occultation and eclipse (of J2 by J3) successively.
- (5) The relative velocity between the 2 involved satellites is too small to allow a good precision on the date of the phenomenon.
- (6) 1985/6/17: total occultation of J2 by J3 from 13^h17^m43^s to 13^h19^m51^s
- (7) 1985/6/24: total occultation of J2 by J3 from 16^h19^m25^s to 16^h21^m35^s
- (8) 1985/7/01: total occultation of J2 by J3 from 19^h19^m30^s to 19^h20^m47^s
- (9) 1985/7/26: total occultation of J2 by J4 from 12^h34^m08^s to 12^h35^m47^s
- (10) 1985/9/04: total eclipse of J2 by J3 from 21^h44^m08^s to 21^h51^m29^s
- (11) 1985/9/12: total occultation of J2 by J3 from 0^h03^m31^s to 0^h16^m24^s
- (12) 1985/11/04: total occultation of J2 by J4 from 2^h52^m21^s to 2^h54^m50^s
- (13) 1985/12/25: total occultation of J1 by J4 from 2^h38^m36^s to 2^h39^m03^s
- (14) 1985/12/29: total occultation of J2 by J1 from 13^h45^m08^s to 13^h45^m30^s
- (15) 1986/1/04: total occultation of J1 by J3 from 19^h36^m14^s to 19^h37^m10^s
- (16) 1986/1/11: total occultation of J2 by J4 from 9^h56^m00^s to 10^h09^m40^s
- (17) 1986/1/26: total occultation of J2 by J3 from 12^h07^m16^s to 12^h08^m54^s
- (18) 1985/9/12: occultation of J2 by J3, the given day is that of the maximum, the beginning of the phenomenon occurring on 1985/9/11, 23^h45^m29^s
- (19) 1986/1/03: occultation of J4 by J2, the given day is that of the maximum, the beginning of the phenomenon occurring on 1986/1/2, 23^h58^m08^s

N.B.: 1) The dates given in the notes are in Ephemeris Time, for G–5 Ephemeris

2) Because the hour of the maximum is the same for the 3 ephemerides (Columns {7}–{9} of Table 1), the minutes may be larger than 60.

velocity between the involved satellites, the greater the precision (generally about 6 s of time).

Since we must use a complete theory to calculate mutual phenomena, much computer time is needed. For the 1979 period we used methods to decrease the computer time (Arlot, 1978). In spite of the technical evolution of the computers, it is always necessary to first determine approximate dates where phenomena are possible. This first determination may cause the omission of some phenomena detectable only with the complete theory.

For the calculations of the dates of the phenomena, we used several ephemerides of the Galilean Satellites: G–5 (Arlot, 1982), E–2 (Lieske, 1980) and SV2 (Vu, 1977; Sampson, 1921). The calculations differ slightly (cf. Table 1). During the 1979 period of phenomena, comparisons were made between these ephemerides and observations (Arlot et al., 1983). For the present predictions

we used the ephemerides of the planets of the JPL DE 102 (Newhall et al., 1983) and the radii of the satellites given by Pioneer 11 (Hollingsworth Smith, 1978). These radii are: for J1, 1840 km; for J2, 1552 km; for J3, 2650 km, and for J4, 2420 km.

Table 1 gives the dates of the predicted phenomena. These dates are given in the TDB time-scale (temps dynamique barycentrique, cf. Connaissance des Temps for 1984). This time-scale is very close to the Ephemeris Time (ET) or to the International Atomic Time plus 32 s (TAI+32 s). One obtains the times of the mutual phenomena in the Universal Time scale (UT) by taking off about 55 s of time from the published dates of Table 1.

The first three columns of Table 1 (Columns {1}–{3}) give the year, the month and the day of the instant of maximum of the considered phenomenon. In Column {4} is given the nature of the phenomenon: 102 means J1 occults J2; 3E4 means that J3

eclipses J4; P means partial phenomenon – nothing is indicated when an eclipse is by the penumbra – A means annular and T means total (dates of the beginning of the totality are given in the notes). The next two columns give, for G–5 ephemeris: first Column {5} the entry in the penumbra of the eclipsed satellite for eclipses, and second, Column {6} the entry in the shadow of the eclipsed satellite or the beginning of the occultation. The next columns give: the date of the maximum of the phenomenon for 3 ephemerides G–5 (Column {7}), E–2 (Column {8}) and SV 2 (Column {10}); the duration of the phenomenon for the same 3 ephemerides G–5 (Column {10}), E–2 (Column {11}) and SV 2 (Column {12}) – for the eclipses, this duration excludes the penumbra – the amplitude of the phenomenon (for the occultations the calculation is made as for the moon¹; for the eclipses, the magnitude is the percentage of the eclipsed satellite in the shadow) for G–5 (Column {13}), E–2 (Column {14}), and SV 2 (Column {15}); the impact parameter (the smallest apparent distance in arcsec between the 2 satellites for an occultation and between the eclipsed satellite and the axis of the umbral cone for an eclipse) for G–5 (Column {16}), E–2 (Column {17}) and SV 2 (Column {18}); the distance of the occulted or eclipsed satellite to the center of Jupiter in jovian radius (Column {19}) and some notes in Column {20}).

In Table 1, mutual phenomena occurring behind Jupiter or in the shadow of Jupiter are not given as phenomena occurring when the apparent distance Jupiter-Sun is less than 10°. The notes indicate:

(1) The phenomena for which the apparent distance Jupiter-Sun is less than 25°.

(2) The grazing phenomena: their amplitude are less than 0.03 for the occultations and less than 2% for the eclipses. Such events are very difficult to observe. Anyway, it can appear that the real amplitude is larger than the predicted one, so the phenomenon is easily observable. As a matter of fact, note that for the eclipse of J 2 by J 3 on July 23, 1985, the calculated amplitude is 11% using SV 2 ephemeris and 0% using G–5 or E–2. On October 5, 1985, for the occultation of J 2 by J 1, the amplitude is 0.239 using SV 2 and 0. with E–2 or G–5 (no event is found). On the contrary, on April 9, 1986, for the occultation of J 1 by J 2, no event is found using SV 2 when an occultation of amplitude 0.190 is found using G–5 or E–2. For such phenomena, the observation will give valuable information.

(3) The phenomena occurring at a distance to the limb of Jupiter less than 0.3 jovian radius or occurring in front of Jupiter (total occultation of J 2 by J 4 on November 4 in 1985, partial occultation of J 1 by J 3 on November 29, 1985, partial occultation of J 1 by J 2 on April 6, 1986 and partial eclipse of J 1 by J 2 (3%) on April 9, 1986).

(4) A special phenomenon: an occultation of J 2 by J 3 just followed by an eclipse of J 2 by J 3.

(5) Very long phenomena: during them the 2 implied satellites have the same apparent speed so that the precision of the calculated dates is not good.

1 for the occultations, the amplitude is calculated as follows: let R_T be the radius of the cone of visibility of the total occultation near the Earth, R_p the radius of the cone of visibility of the partial occultation near the Earth and D the distance between the Earth and the common axis of these cones. The amplitude is $\frac{R_p - D}{R_p - R_T}$ (Danjon, 1959, p. 320)

Table 2. Distribution of the different mutual events during the 1985–86 occurrence

	NUMBER OF OCCULTATIONS							NUMBER OF ECLIPSES										
	1985 feb. jul.	aug.	sept.	oct.	nov.	dec.	1986 jan. apr.	SUMS	1985 feb. jul.	aug.	sept.	oct.	nov.	dec.	1986 jan. apr.	SUMS		
I/I	0	0	7	8	8	8	5	0	37	31	0	3	9	10	8	1	0	0
II/I	2	0	0	0	0	1	9	11	23	51	0	0	0	8	8	9	9	17
I/III	3	5	5	0	0	4	4	0	21	18	0	3	8	4	1	0	2	0
III/I	10	0	0	0	4	5	3	0	22	18	0	0	2	5	8	2	1	0
I/IV	4	2	2	0	0	2	1	0	11	5	0	1	2	2	0	0	0	0
IV/I	7	2	0	0	2	4	0	0	15	38	7	0	2	0	3	2	0	0
II/III	2	0	0	0	0	4	1	3	10	7	1	0	0	0	1	0	0	5
III/II	10	4	8	0	0	0	3	1	26	64	3	5	4	0	4	4	1	0
II/IV	2	1	0	0	1	0	1	0	5	5	1	0	1	2	1	0	0	0
IV/II	2	0	1	0	2	0	2	0	7	21	4	0	0	2	1	1	0	0
III/IV	2	1	0	0	0	0	1	0	4	4	0	0	2	1	0	1	0	0
IV/III	2	1	0	0	0	1	1	0	5	17	4	0	1	1	1	1	0	0
SUMS:	46	16	22	8	17	29	32	15	186	175	5	15	31	37	35	17	13	22
									361									

The use of the radius deduced from the observations made by Voyager – J 1: 1816 km, J 2: 1563 km, J 3: 2638 km, J 4: 2410 km (Morrison, 1983) – changes very slightly the results (anyway the dates of the maximum remain unchanged). The greatest difference in duration is about 20s for slow phenomena, and most of the dates of the beginning of the phenomena change only by less than 10s.

Conditions for the observation of the mutual phenomena in 1985–1986

How are the observable events distributed in 1985–86? It would be nice to have the same number of events for each pair of satellites. However, since the period of each satellite is different, the faster satellites participate in more events than the slower ones. Table 2 gives the number of phenomena for each pair of satellites each month. It is shown that the pair 1–2 participates in many more events than the other pairs, and that the pair 3–4 participates in very few events. The period between maximum events is not the same for each pair because of the small, distinct, inclination of each satellite. Note that the radii of the satellites also determine the number of events.

Table 3 gives the number of events as a function of the duration (the transit of eclipsed satellites in the penumbra is not taken into account in the calculation of the duration). The longest event is about 2 h long and 17 events are about one hour long.

Table 4 gives the number of events (grazing or not) observable in several observatories. We selected phenomena occurring, first, when Jupiter is more than 5° above the horizon and second, when Jupiter is more than 15° above the horizon. In all cases, the Sun is less than 5° below the horizon. It is easy to understand that the best observational conditions are obtained for the observatories whose latitudes are close to the declination of Jupiter (about –19° to

Table 3. Numbers of mutual events as a function of their durations

DURATIONS	NUMBER OF		SUMS
	OCCULTATIONS	ECLIPSES	
GRAZING EVENTS	32	81	113
< 100 ^s	0	1	1
100 ^s TO 299 ^s	57	46	103
300 ^s TO 599 ^s	43	20	63
600 ^s TO 999 ^s	26	10	36
1000 ^s TO 1499 ^s	9	6	15
1500 ^s TO 2499 ^s	8	4	12
2500 ^s TO 4999 ^s	11	6	17
5000 ^s TO 7000 ^s	1	0	1
SUMS :	186	174	361

– 14° during the period of the mutual phenomena). If someone wishes to have the list of phenomena observable in a special observatory, please contact the author.

The instruments to be used for the observation of the mutual phenomena

Since a large number of astronomers made observations of mutual phenomena during the 1973–74 and 1979–80 periods, most of the observers are aware of the main instruments which can be used for these observations. Everyone knows that the magnitude of the galilean satellites allows their observations with telescopes of small apertures. The observation of nongrazing phenomena is possible even with a sky of bad transparency (in cities for example). On the contrary, to catch grazing phenomena, very good conditions are needed.

To be useful, the observation of a mutual phenomenon may give at least a date in UT (precise within half a second of time) of the instant of the maximum drop of light during the phenomenon (which corresponds to a minimum intensity on the light curve). Anyway, other elements of the observation are of great interest: – the entire light-curve will be useful to improve the accuracy of the time of the instant of the drop of light and necessary to calculate the duration of the phenomenon; – an evaluation of the magnitude drop will be necessary to determine the amplitude of the phenomenon. All these data (time of the minimum intensity light, duration and amplitude) are useful to improve the ephemerides.

Since the sky background may vary (mainly during twilight), it is necessary to record it several times during each observation.

For these photometric observations, two possibilities of recording are available. First, the old classical single channel photoelectric photometer with analog or digital output may be used. In that case, except for the eclipses where the eclipsing satellite is far from the eclipsed one, both satellites involved must be put in the diaphragm. Second, two-dimensional photometers (for example Vidicon-tubes) may be used for the observations. In that case, there is no problem of diaphragm since the images are studied numerically after the observations. The large number of data constrain to record the phenomenon analogically and to digitize the images after the observation. This method has been already used successfully for the observation of mutual phenomena (Mosher et al., 1975; Arlot et al., 1982) or for other observations of the same type (Arlot et al., 1984).

Table 4. Numbers of observable phenomena (not grazing) for several observatories

OBSERVATORIES		NUMBER OF OBSERVABLE PHENOMENA (NOT GRAZING)	
NAMES	LATITUDES	JUPITER IS MORE THAN 5 DEG. ABOVE THE HORIZON	JUPITER IS MORE THAN 15 DEG. ABOVE THE HORIZON
LA SILLA (CHILI)	–29°	66	57
SUTHERLAND (SOUTH AFRICA)	–32°	69	56
JOHANNESBURG (SOUTH AFRICA)	–26°	62	56
CAPE TOWN (SOUTH AFRICA)	–34°	67	54
RIO DE JANEIRO (BRASIL)	–23°	60	51
MELBOURNE (AUSTRALIA)	–38°	50	46
MAUNA KEA (HAWAI ISL.)	+20°	54	44
WISE, OBS. (ISRAËL)	+30°	55	45
HELWAN (EGYPT)	+30°	51	44
GRANADA (SPAIN)	+37°	54	43
MADRID (SPAIN)	+40°	52	43
KODAIKANAL (INDIA)	+10°	51	43
MC DONALD (USA)	+31°	48	42
HYDERABAD (INDIA)	+17°	49	40
CATANIA (ITALY)	+37°	49	41
PIC-DU-MIDI (FRANCE)	+43°	48	41
BARCELONA (SPAIN)	+41°	44	41
MOUNT PALOMAR (USA)	+33°	50	39
LYON (FRANCE)	+46°	47	40
BORDEAUX (FRANCE)	+45°	47	40
CERGA/NICE (FRANCE)	+44°	45	40
MC CORMICK (USA)	+38°	48	38
FLAGSTAFF (USA)	+35°	47	38
OBS. HTE PROVENCE (FRANCE)	+44°	45	39
SHANGHAI (CHINA)	+31°	42	38
DODAIRA (JAPAN)	+36°	51	37
ROMA (ITALY)	+42°	47	38
PURPLE MOUNTAIN (CHINA)	+32°	42	37
USNO (USA)	+39°	47	36
MITAKA (JAPAN)	+36°	46	34
PARIS-MEUDON (FRANCE)	+48°	45	35
NAINITAL (INDIA)	+29°	44	34
PEKING (CHINA)	+40°	41	34
DRESDEN (GERMANY)	+51°	52	30
PULKOVO (URSS)	+60°	32	0
OULU (FINLAND)	+65°	20	0

Another method of observation is the visual one. With this method, it is not possible to measure accurately the magnitude drop. Anyway, the time of the instant of maximum of the phenomena may be determined by reference to a nearby satellite not affected by the phenomenon. Trained observers (especially amateur astronomers) are able to provide useful data: this has been demonstrated during the 1979–80 period (Arlot et al., 1982).

The significance of the observations of mutual phenomena

These observations should be made in order to improve our knowledge of the Galilean Satellites. The former campaigns of observations produced a large set of data which has been used for several purposes: for the improvement of the ephemerides (mid-times of the events for the longitude corrections and magnitude drops for the latitude corrections), for the determination of the radius (duration of the phenomena) and for the evaluation of the albedos (magnitude drops and shapes of the light-curves).

A large number of publications were made after the last two periods of occurrence of mutual phenomena which increased our knowledge on the subject. But now, is it always worthwhile to make such observations? Spacecrafts Pioneer and Voyager gave valuable information about the Galilean Satellites. As a matter of

fact, some quantities which were not known have been determined with a high accuracy by these spacecrafts: for instance radius and surface features. On the other hand, the albedos are not yet well known and the asymmetry of some light-curves of mutual phenomena still remains unexplained. About the ephemerides, more observations are necessary. Let us notice that ephemerides may be improved using observations of eclipses by Jupiter or photographic observations. Moreover the observations of mutual phenomena are more precise than other types of observations. Photographic observations, however, may be made at anytime whereas mutual phenomena are fairly rare. But the observations of mutual phenomena are very different from the other types of observations. They provide information not otherwise available. Moreover, it is easier to eliminate the systematic observational errors when several types of observations are available. It is valuable to mix different types of observations of Galilean Satellites, when calculating the constants of their motions.

Conclusion

Observations of the mutual phenomena of the Galilean Satellites are always of great interest and we encourage observers to make them; note that they do not occur each year (if someone needs more information about those observations, please contact the author). In the past, everyone making such observations provided interesting results, but much more may be done and we look forward to much observational data in 1985 and 1986.

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